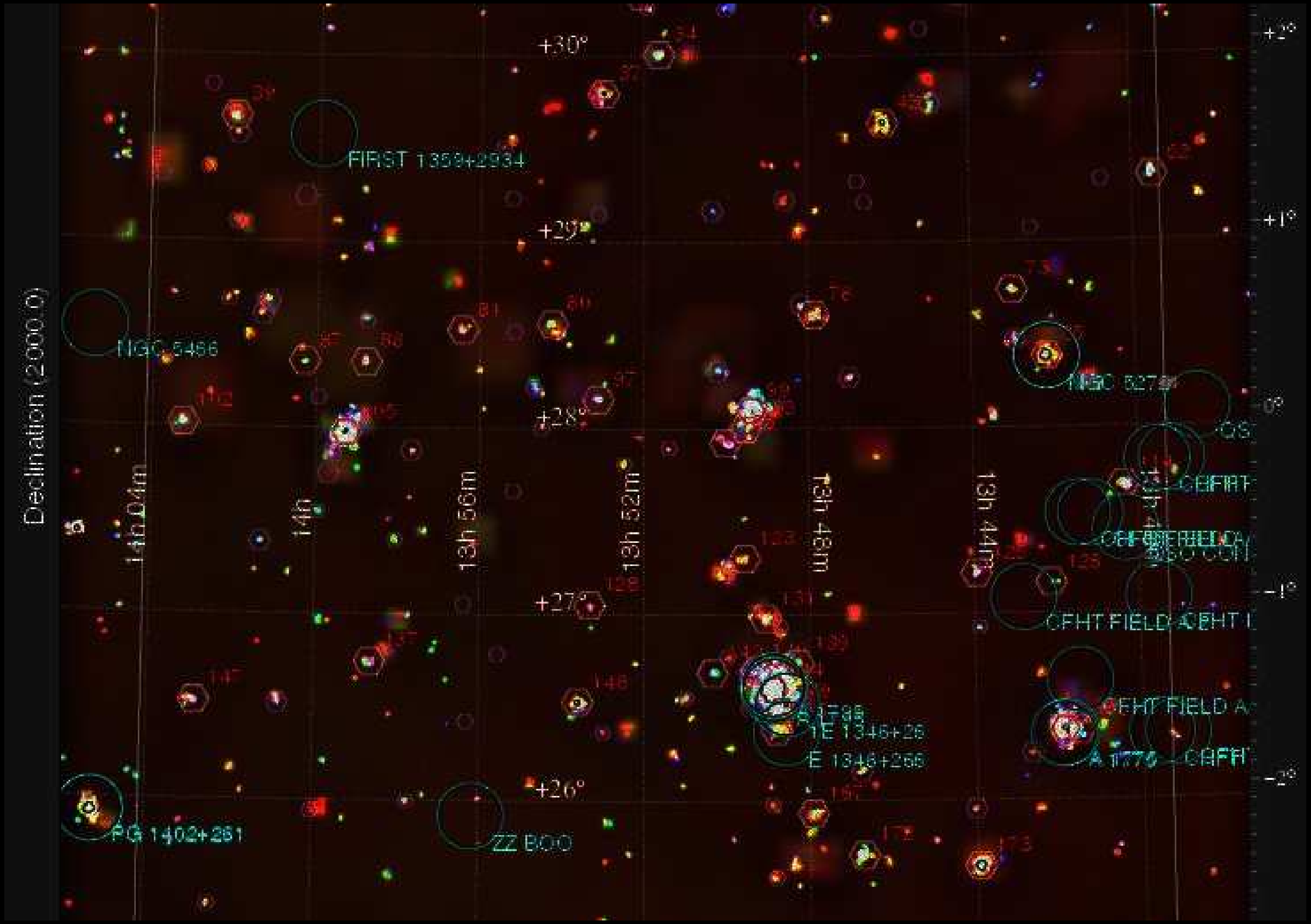
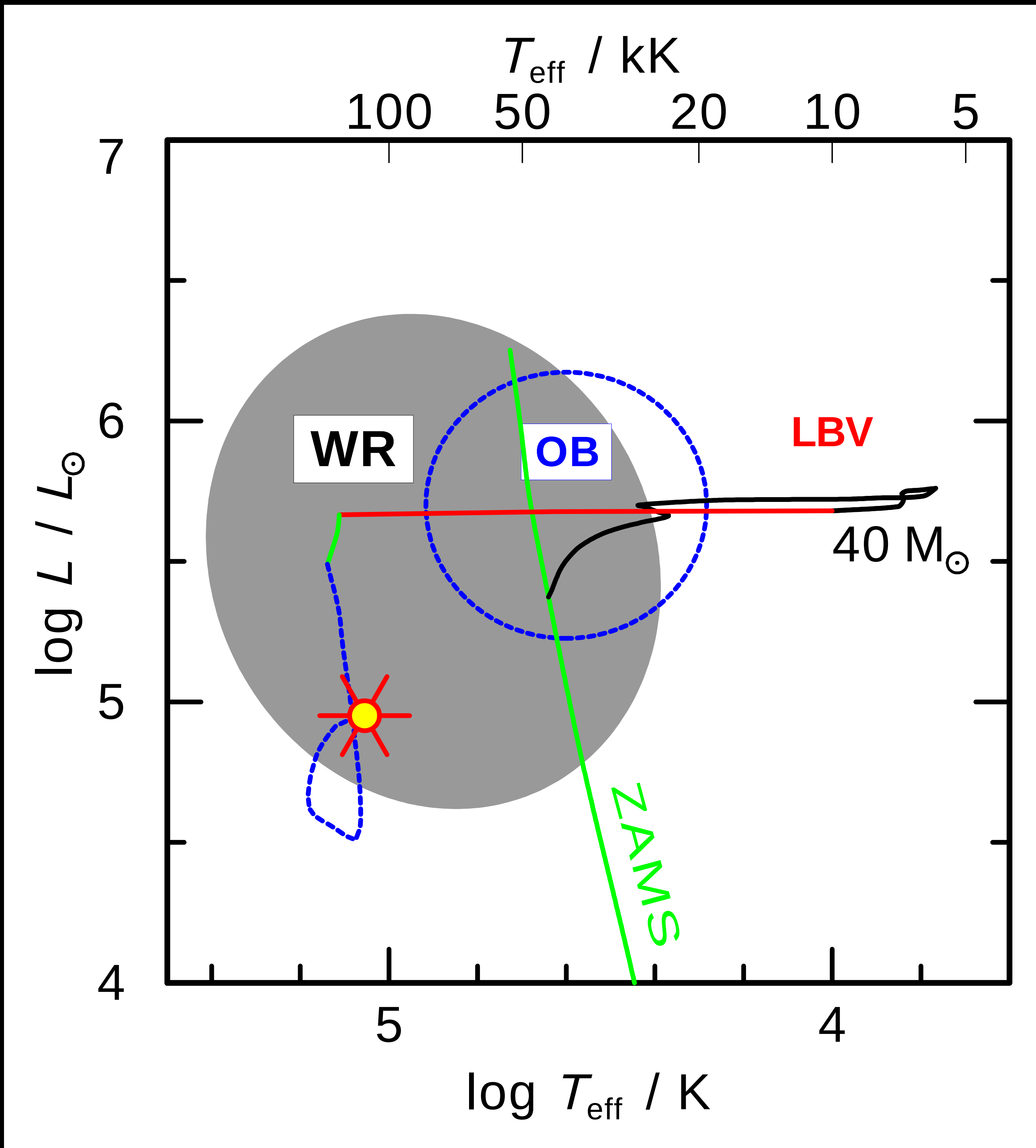
X-rays from massive stars and clusters



Lidia Oskinova
Universität Potsdam

eROSITA @ AIP 2014

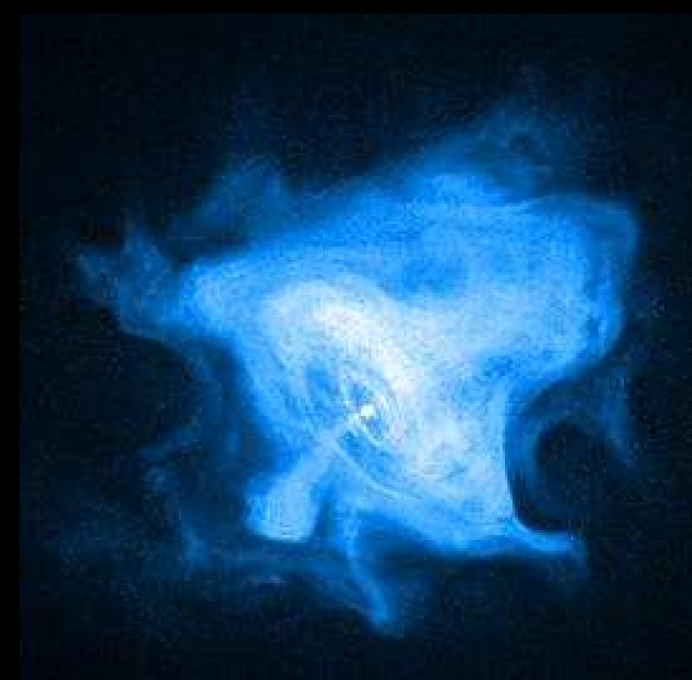
The evolution of (very) massive stars



Mass removal by wind drives the evolution

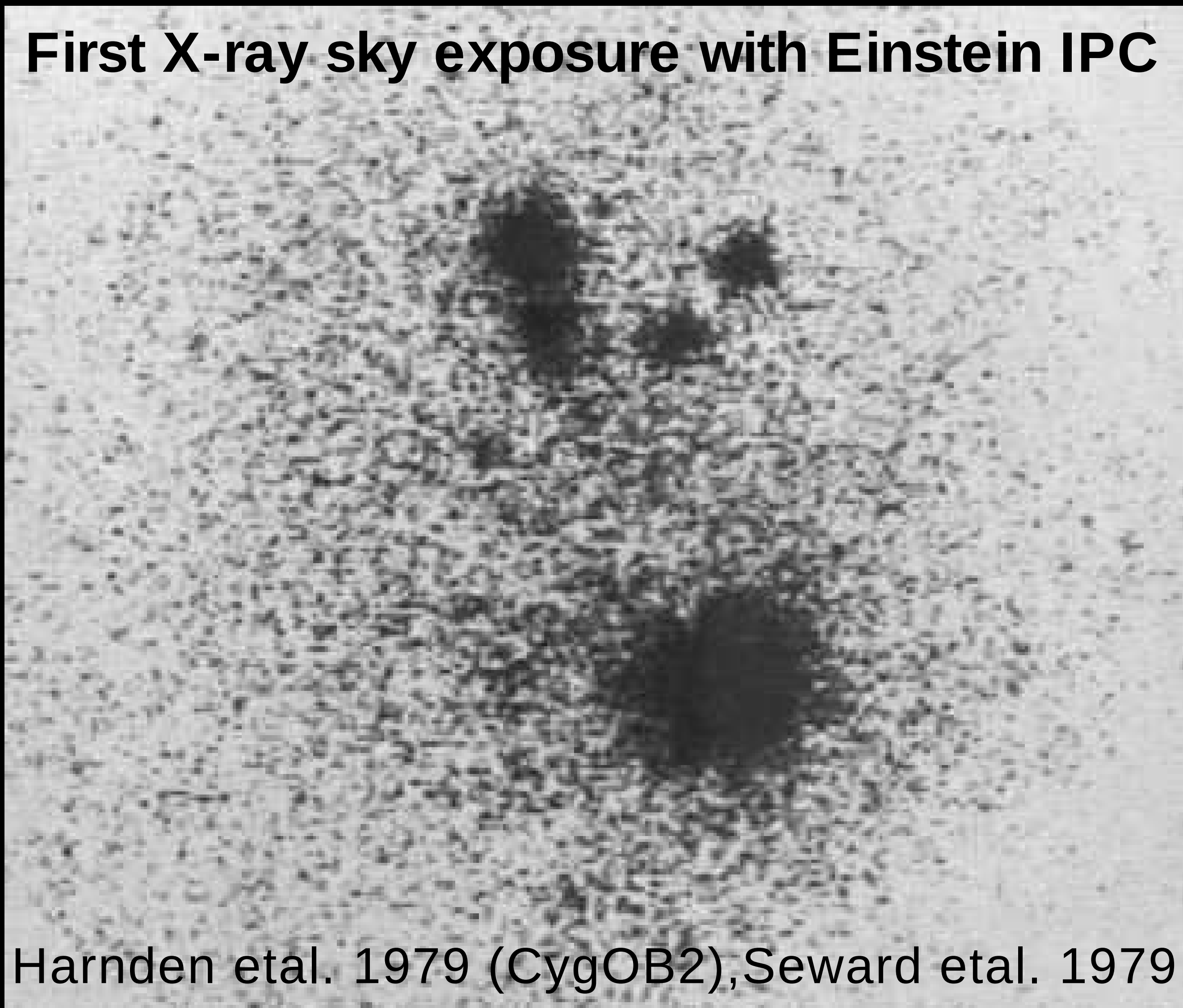
- OB → (short LBV) → WR
- Winds are getting dense and more enriched
- Wolf-Rayet (WR)

WN → WC → WO → SN



Credit: NASA/CXC/SAO/F.Seward et al

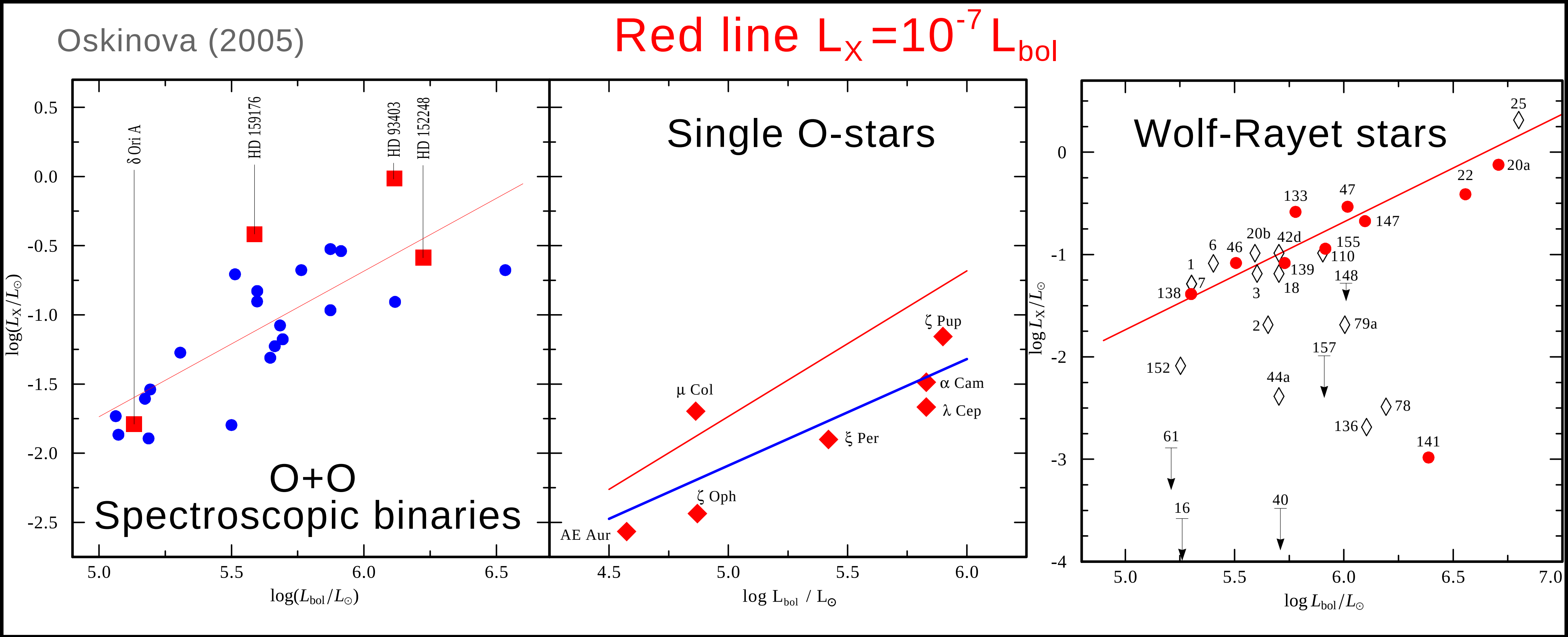
Massive stars emit X-rays (Einstein observatory 1978)



- X-rays were predicted by Hoare (1975), Cassinelly & Olson (1978)
- Seminal study Pallavicini etal. (1981): O and early B stars → Correlation $L_X = 10^{-7} L_{bol}$ but no correlation with $v \sin i$
- The spectra are soft, peculiar stars have harder spectra.

- Important confirmation: RASS (Berghoefer etal. 1997)
- Chandra and XMM-Newton observations (e.g. Moffat+ 2002, Oskinova 2005, Sana+ 2006, Wolk+ 2008, Townsley+ 2009, Gagne+ 2011)

Examples of $L_X = 10^{-7} L_{bol}$ correlation



- The correlation is approximate
- It holds for O and early B stars, single and binaries
- Wolf-Rayet stars do not follow this correlation
- Stars later than B1-spectral type do not follow this correlation

WHY massive stars emit X-rays ?

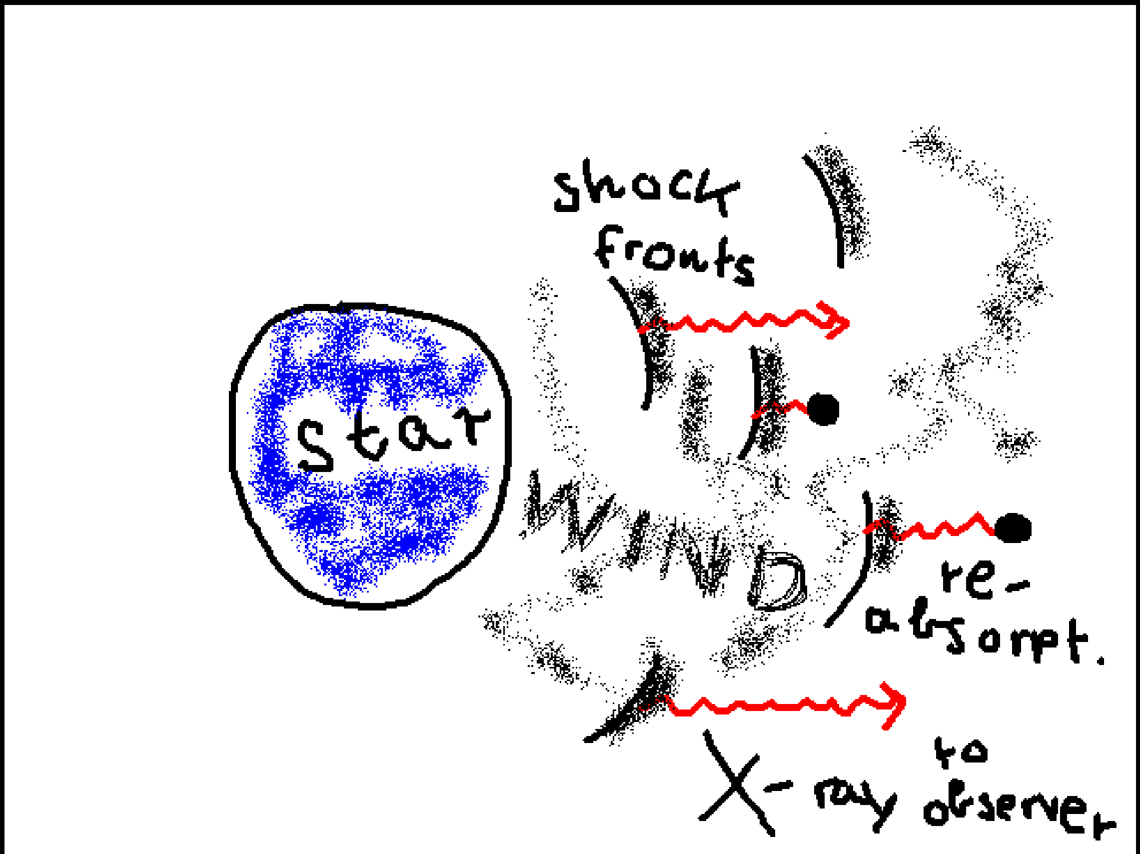
Two leading theories

- Shocked stellar winds
- Hot coronae associated with surface magnetic field

Succesful model shall explain the $L_X = 10^{-7} L_{bol}$ correlation

Radiatively driven stellar winds are intrinsically unstable

Lucy & Solomon (1970), Radiative hydrodinamic X-ray: Feldmeier etal (1997)



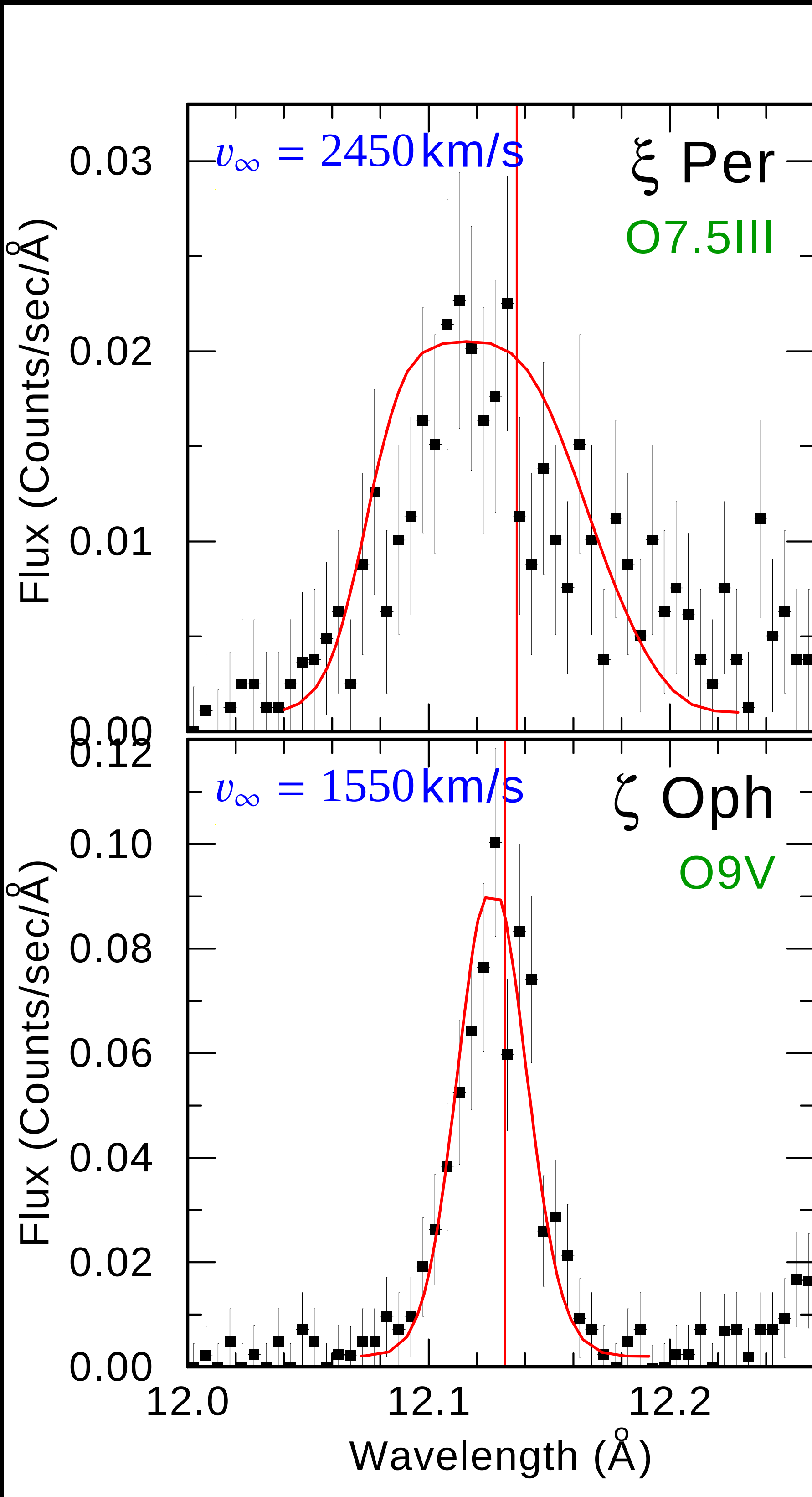
Shocks

Heating

X-Rays

Shocks also result from:

- **Collision of streams in magnetically confined wind**
- **Collision of winds in binaries**

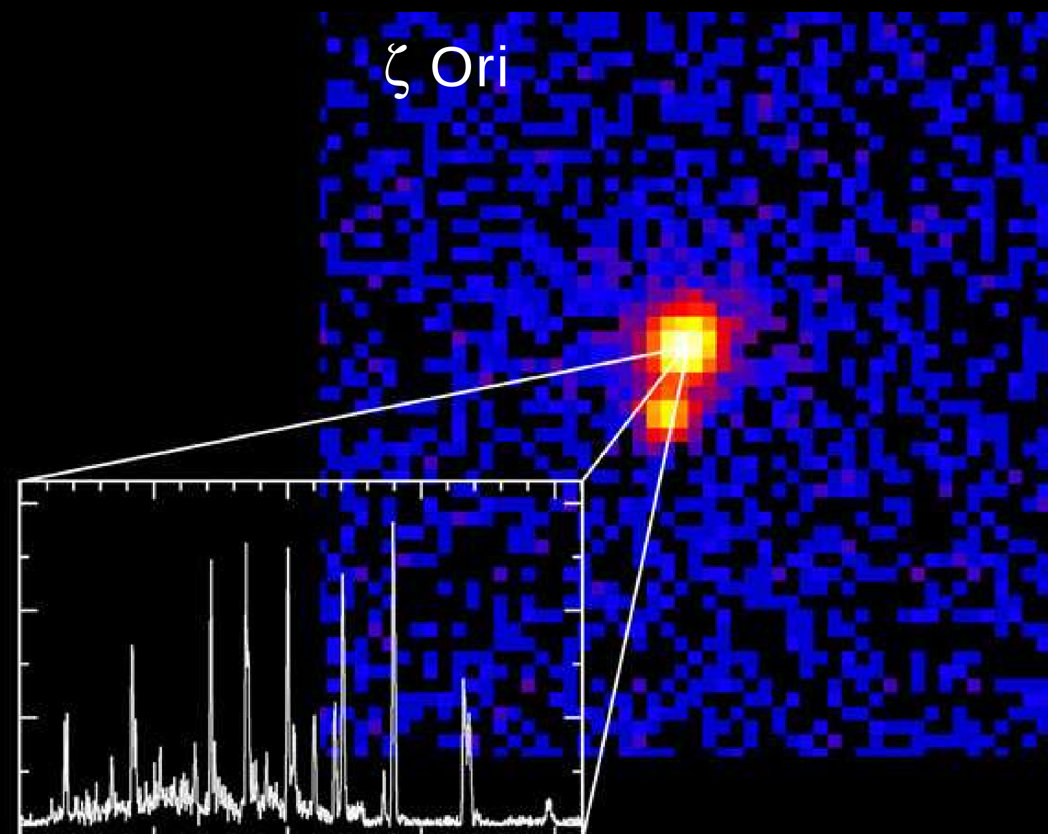


Wind shock model explains:

broad and blueshifted X-ray emission lines

(Cassinelli etal 2001, Kramer etal 2003, Oskinova etal 2006)

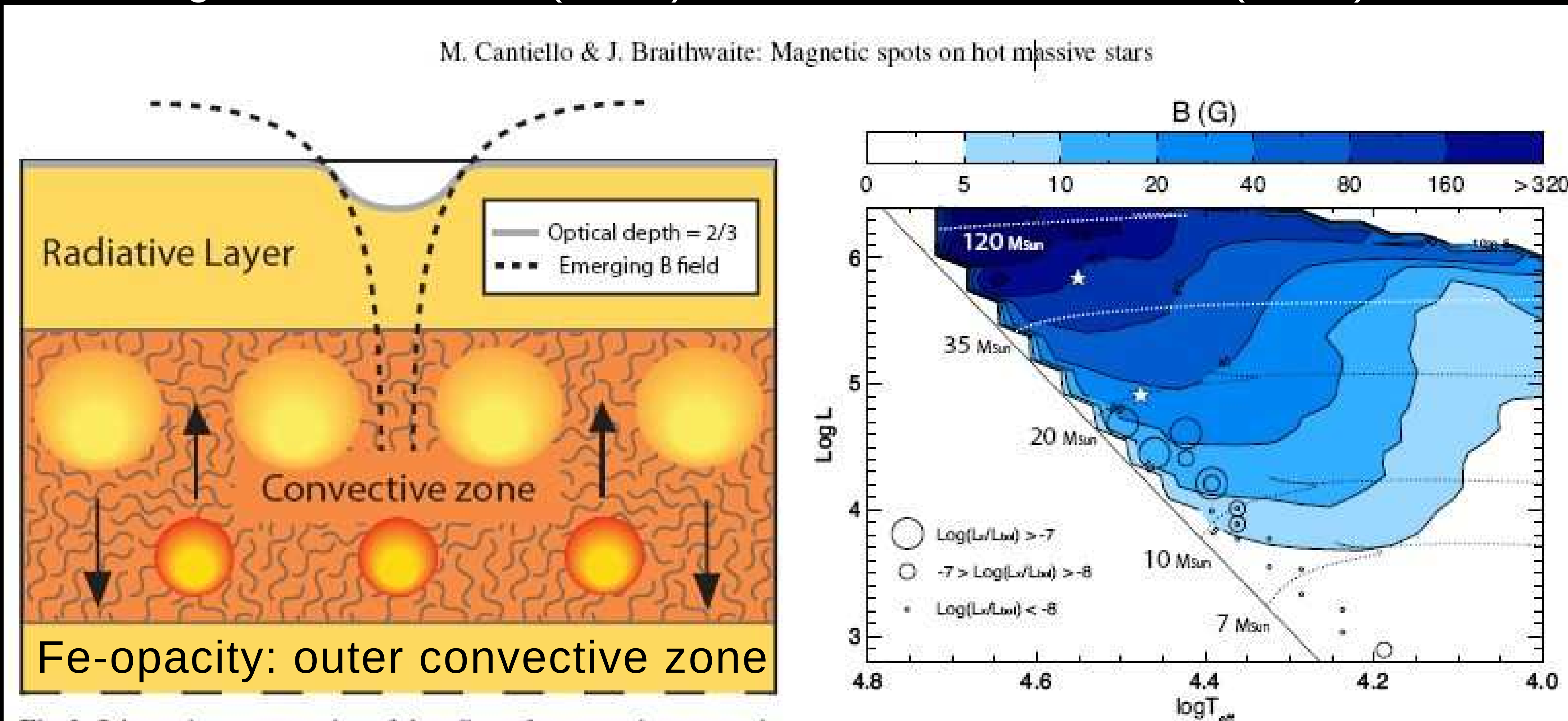
The model does not predict the correlation between L_X and L_{bol}



NASA/EIT/W.Waldron, J.Cassinelli

Surface magnetic field

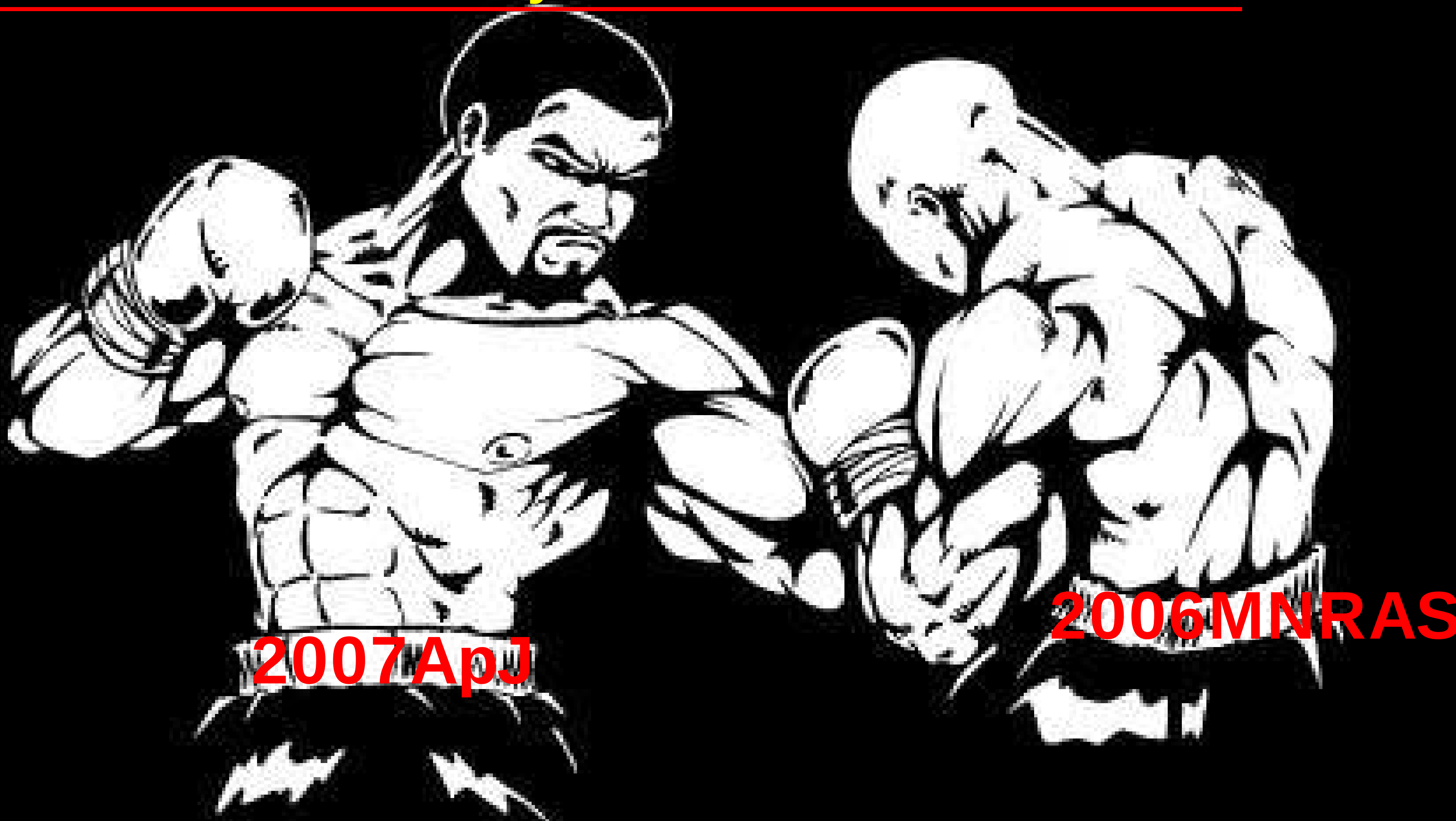
MacGregor & Cassinelli (2005), Cantiello & Braithwaite (2011)



The model:

- Explains why the hottest X-ray plasma is close to the photosphere
- It predicts the dependence of **B**-field strength on $\mathbf{L_{bol}}$
- It may explain $\mathbf{L_X = 10^{-7} L_{bol}}$

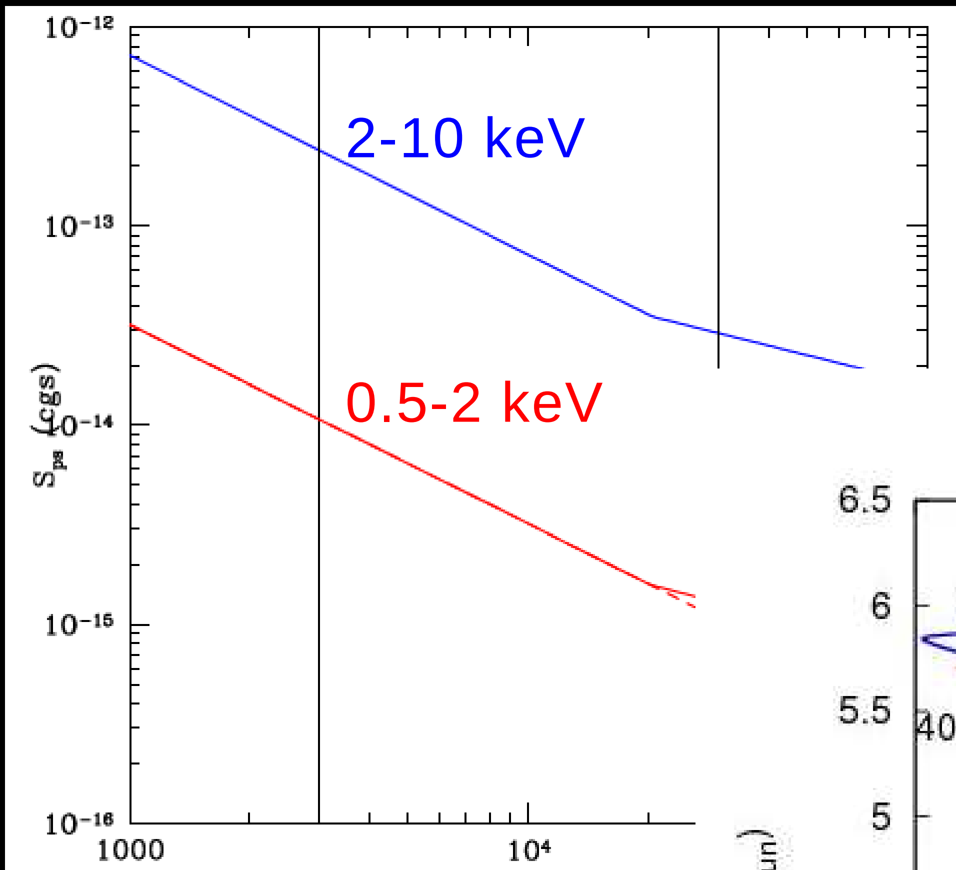
Two rival models of X-ray emission in massive stars



New data are needed to decide among the models:
- High-resolution spectra (XMM-Newton and Chandra)
- Comprehensive sample of stars across the HRD (**eROSITA**)

eROSITA and the evolution of massive stars

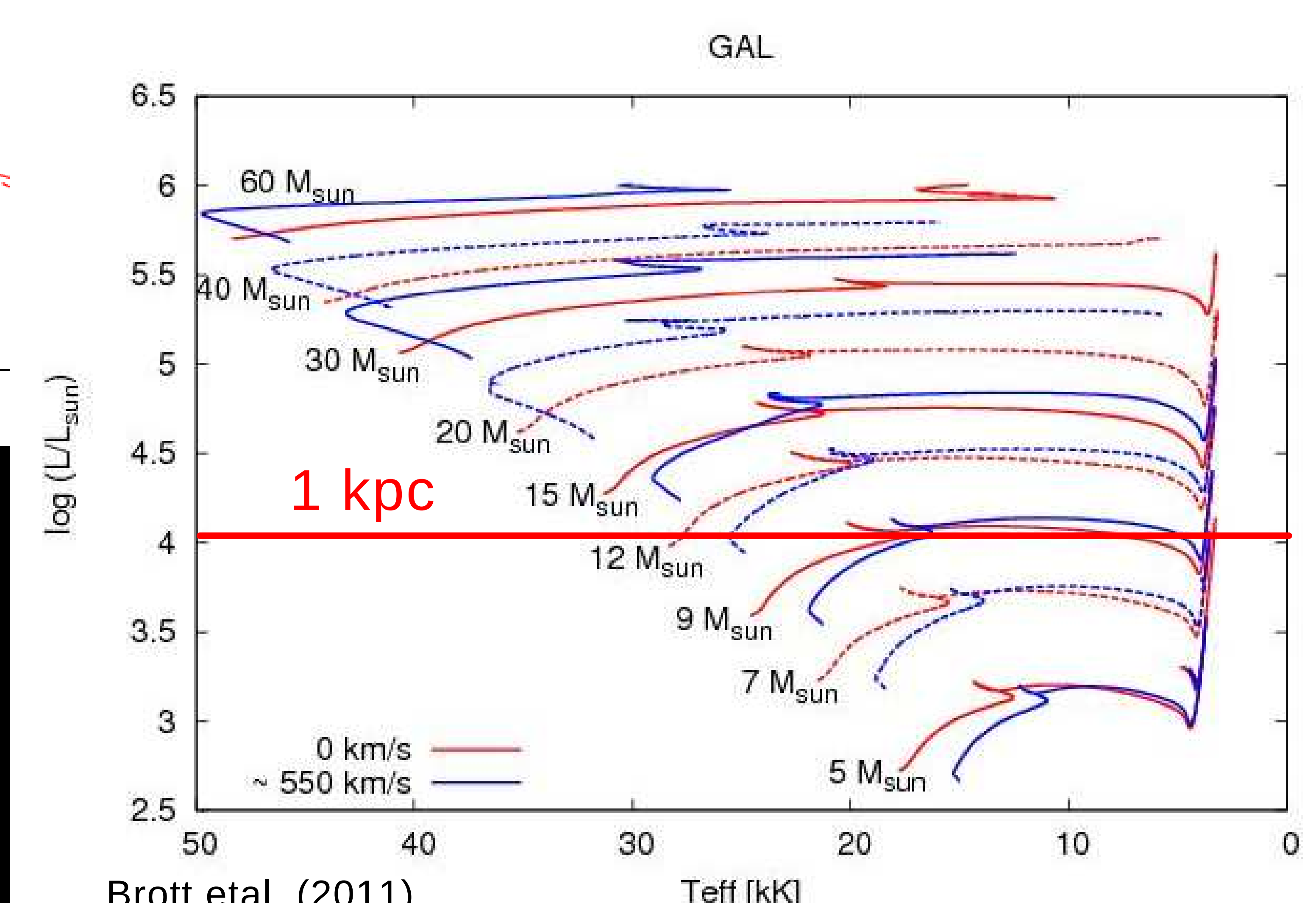
5σ point-source sensitivity vs exp.time



Predehl etal. (2010)

Complete sample within 2 kpc

HRD of massive stars



Brott etal. (2011)

$N_H < 5\ 10^{21}\ cm^{-2}$

all earlier B1

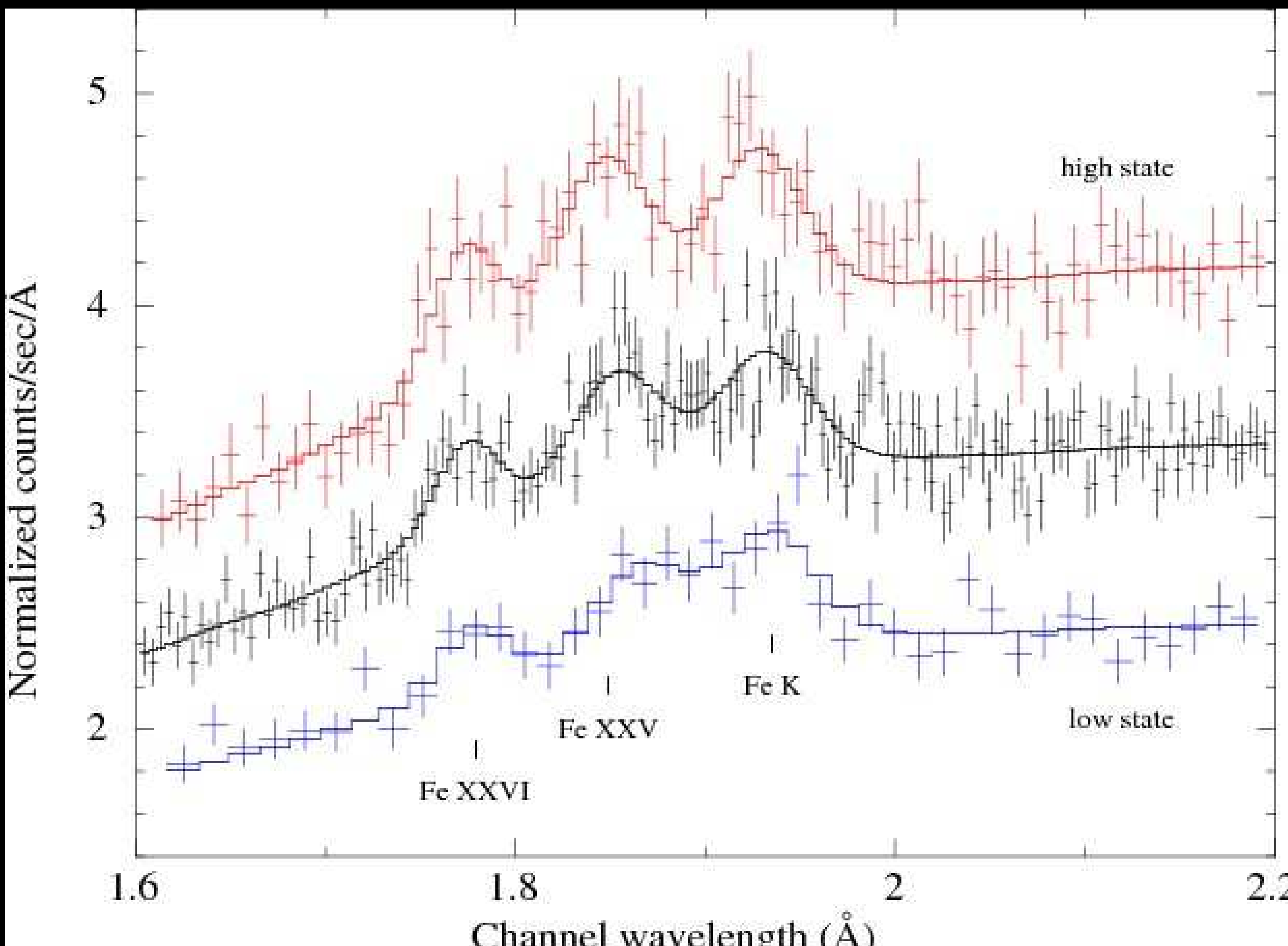
7% magnetic

70% binaries

Look for the surprises in the hard band

eROSITA survey not only 2 dex deeper than RASS.
All-sky survey in hard X-ray band

Variability in Fe complex γ Cas



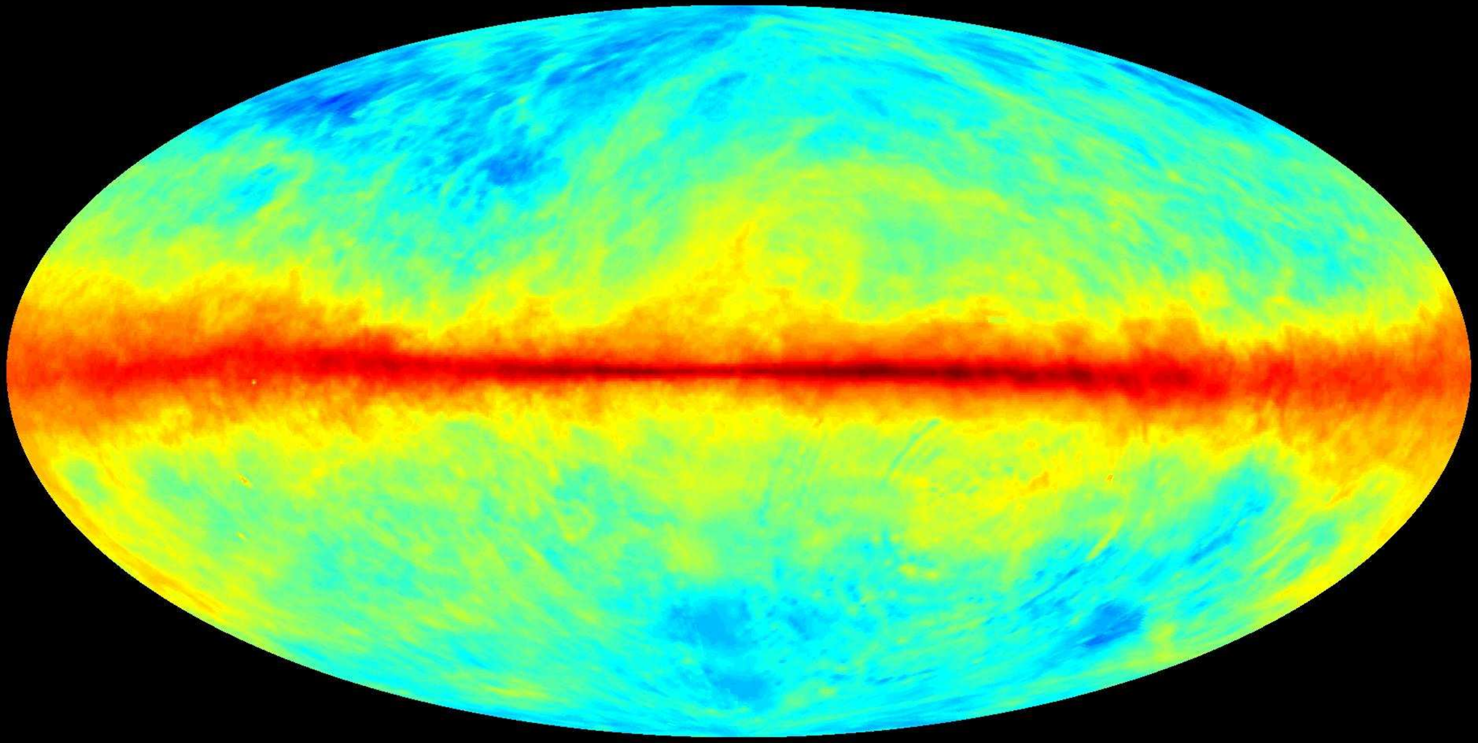
Lopes de Oliveira etal (2010)

- Magnetic stars and colliding wind binaries
- Some massive stars are hard in X-rays, but physical mechanisms are not known e.g:
  γ Cas type objects
  WR46, WR48a
  τ Sco analogs
  θ^2 Ori A stars
  ...

Galactic plane absorption

Massive stars concentrate in galactic plane

Leiden-Dwingeloo 21 cm at δ>-30; Dickey-Lockman 21 cm at δ>-30



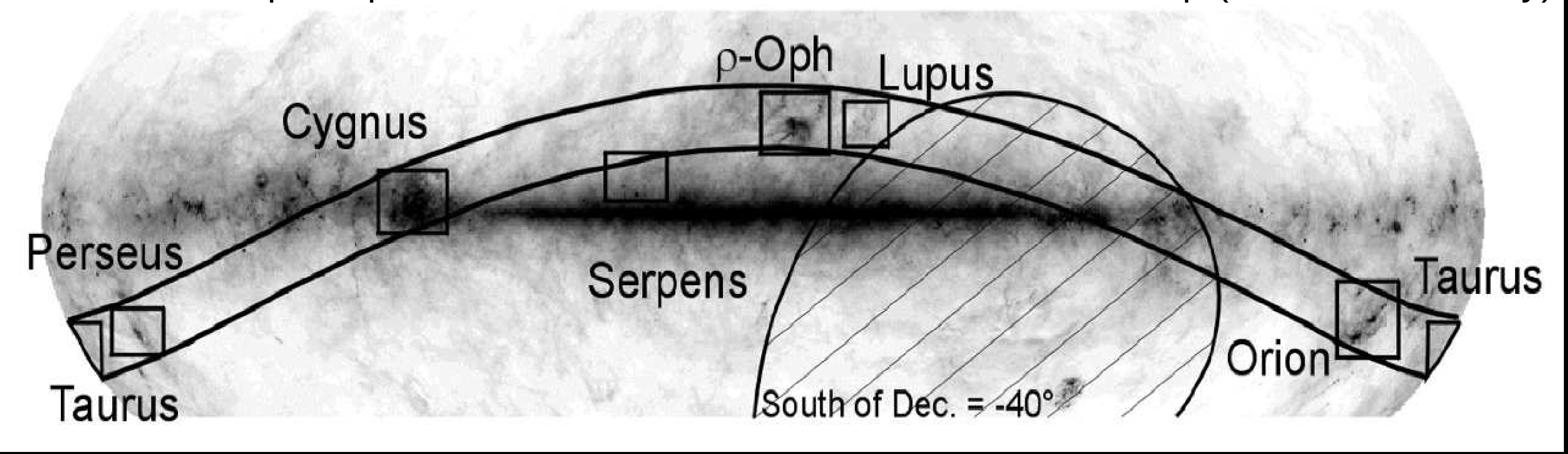
Even in the most obscured regions, eROSITA will reach F_X=6 10^{-13} erg s^{-1} cm^{-2}. Sensitive to detect 'normal' massive star within 1.5 kpc.

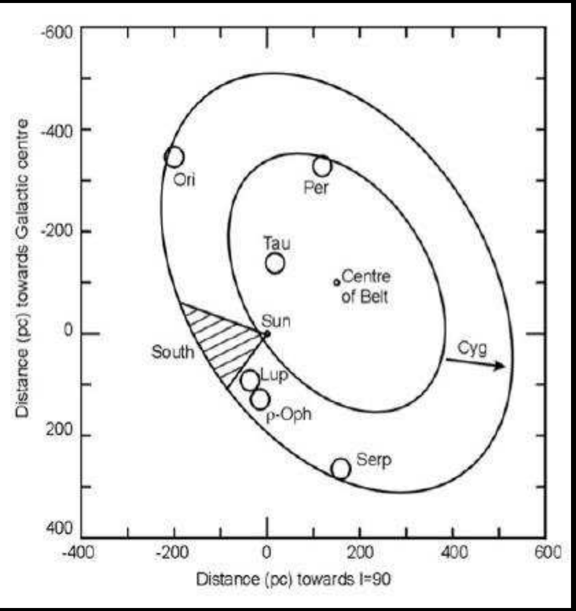
Important: X-ray spectra and hardness ratios

Local Star Formation: Gould's Belt

- Most star formation within 0.5 kpc lies in **Gould's Belt.**
- Ring centred on a point 200 pc from the Sun and tilted at 20 degrees to the Galactic Plane.

Gould's Belt superimposed on to an IRAS 100 micron emission map (JCMT mm survey)

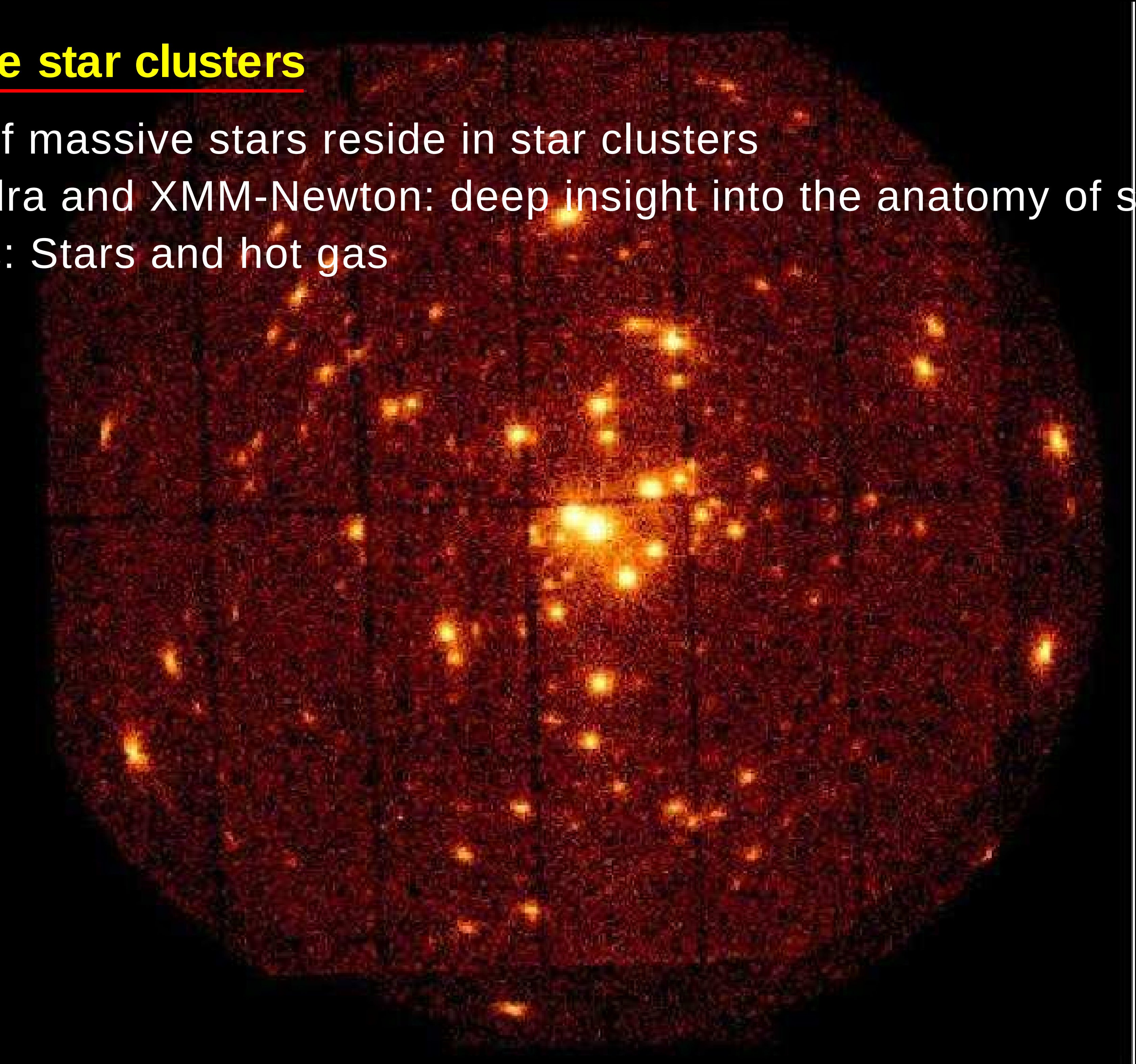




- eROSITA will provide a very significant high-energy picture of solar neighbourhod
- eROSITA combined with IR and mm surveys: unravel **the local star formation history**

Massive star clusters

- 70% of massive stars reside in star clusters
- Chandra and XMM-Newton: deep insight into the anatomy of star clusters: Stars and hot gas



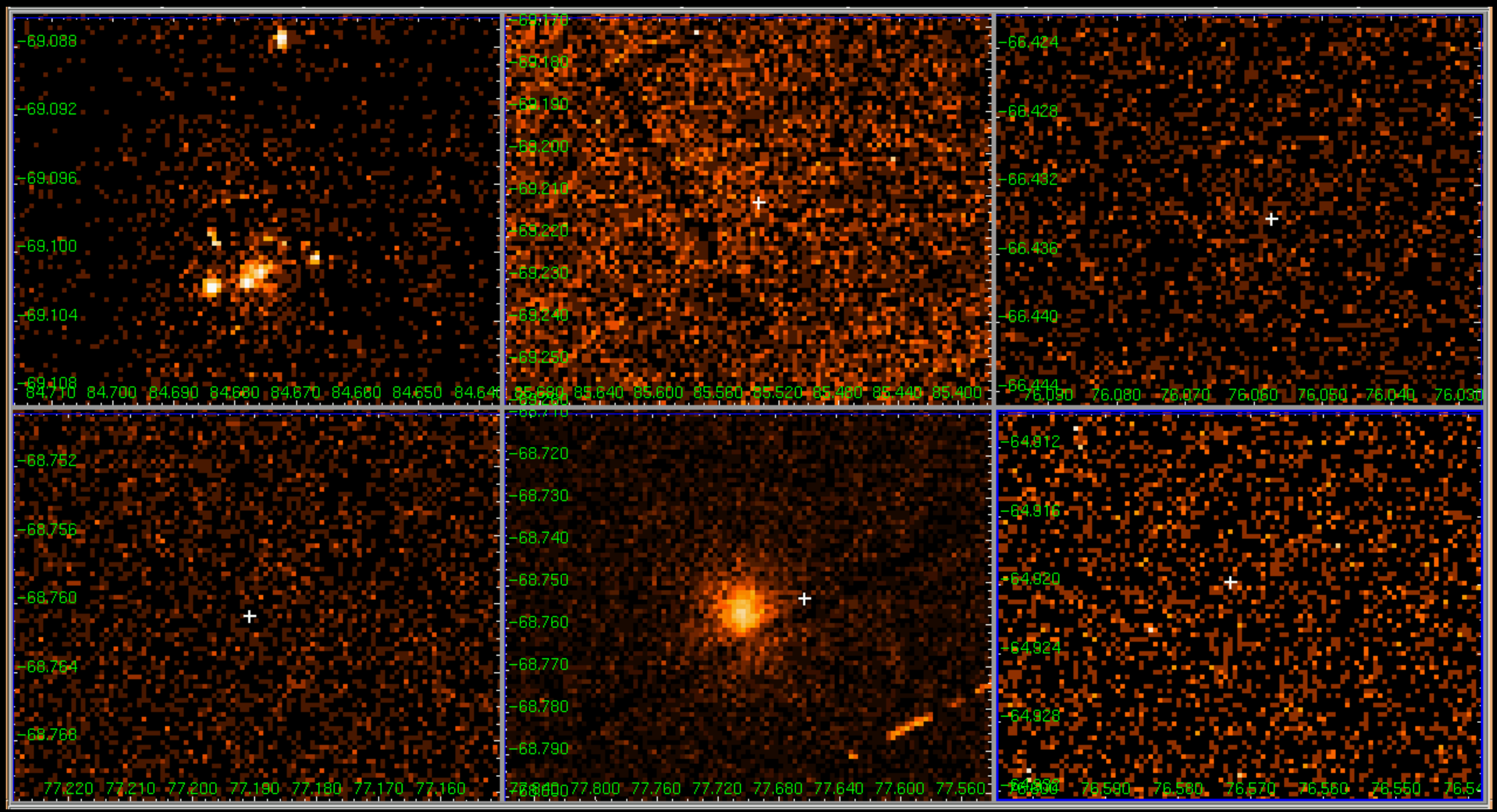
Open stellar cluster: sigma Orionis

Image courtesy of R. Pallavicini, Osservatorio Astronomico di Palermo, Italy

European Space Agency

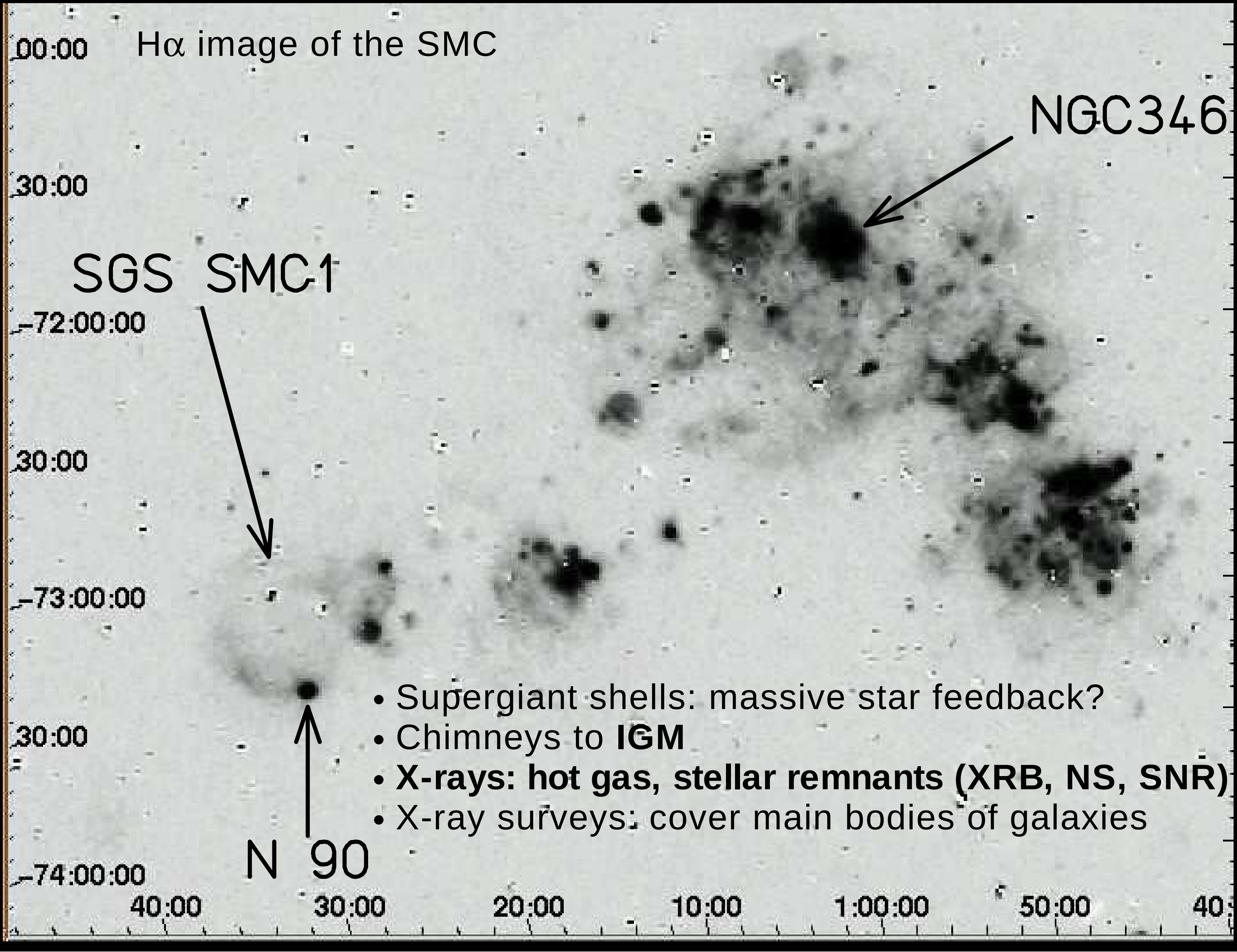
Sample of massive star clusters in the LMC

- LMC, SMC important sample of clusters differemt M and Age
- There is evolution of X-ray emission from clusters
- Well studied in optical/IR: cluster parameters are known



eROSITA: variability, transients, bubbles, XRBs location

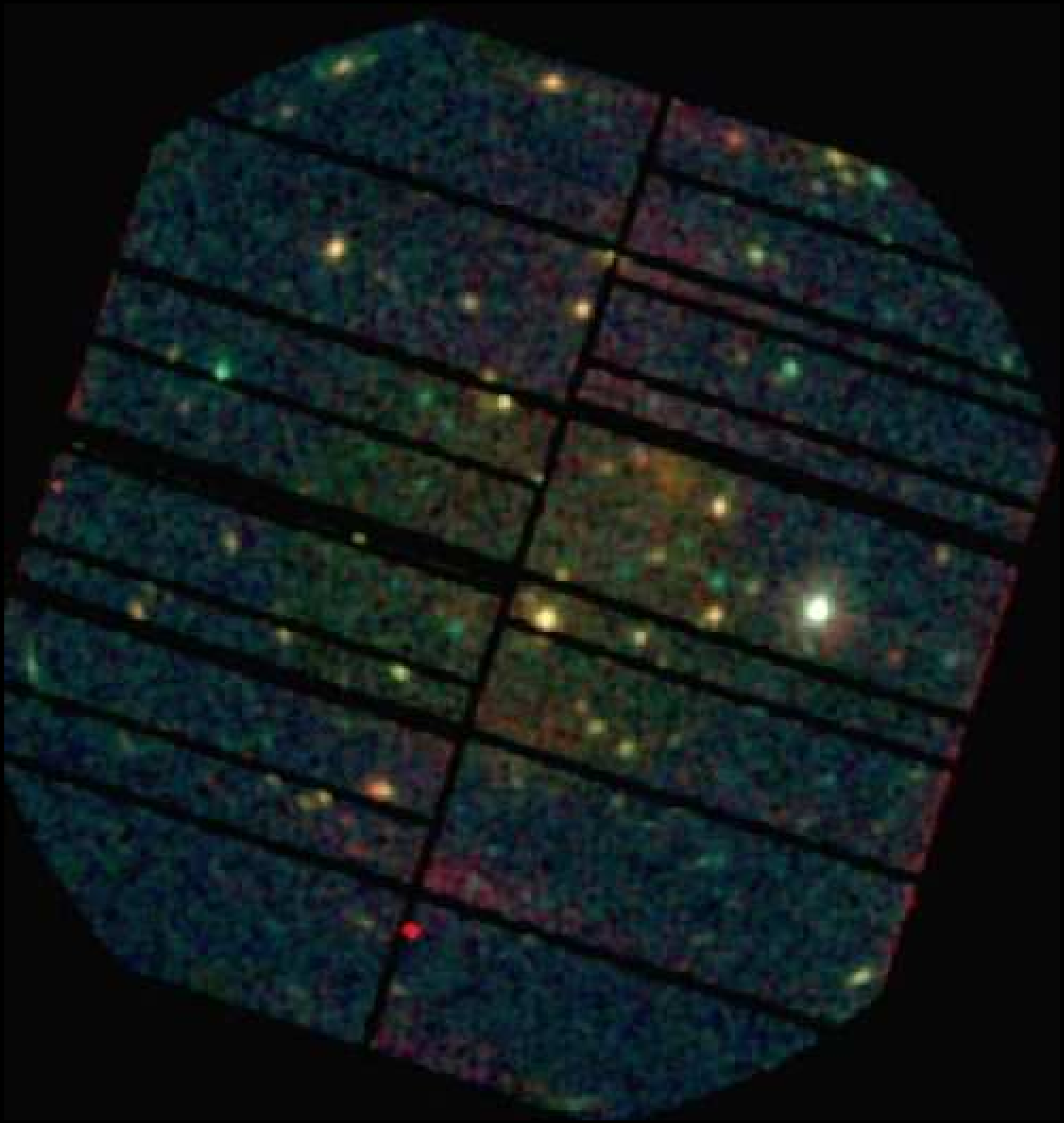
Massive stars and clusters: structuring galaxies



- Supergiant shells: massive star feedback?
- Chimneys to **IGM**
- **X-rays: hot gas, stellar remnants (XRB, NS, SNR)**
- X-ray surveys: cover main bodies of galaxies

Part of the SMC Wing in X-rays

EPIC PN 3 colour



Henault-Brunet etal. (2011), Oskinova etal (2012)

The periferal part of the SMC Wing: low ρ and Z

Star formation: new XRBs are detected

eROSITA: galactic suburbs - starformation in province

LMC: 20ks survey. Excellent to study the SMC Wing, the Bar, the galaxy interactions

XRB with $L_X > 10^{32}$ erg s^{-1} will be identified

Diffuse X-ray emission: heated gas

Optical/IR/radio follow up are important